

Early-type galaxies at large galactocentric radii - I. Stellar kinematics and photometric properties

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ABSTRACT

We present the results of a combined analysis of the kinematic and photometric properties at large galactocentric radii of a sample of 14 low-luminosity early-type galaxies in the Fornax and Virgo clusters. From Gemini South GMOS long-slit spectroscopic data we measure radial profiles of the kinematic parameters v_{rot} , σ , h_3 , and h_4 out to $\sim 1\text{--}3$ effective radii. Multi-band imaging data from the HST/ACS are employed to evaluate surface brightness profiles and isophotal shape parameters of ellipticity, position angle and discyness/boxiness. The galaxies are found to host a cold and old stellar component which extend to the largest observed radii and that is the dominant source of their dynamical support. The prevalence of discy-shaped isophotes and the radial variation of their ellipticity are signatures of a gradual gas dissipation. An early star-forming collapse appears to be the main mechanism acting in the formation of these objects. Major mergers are unlikely to have occurred in these galaxies. We can not rule out a minor merging origin for these galaxies, but a comparison of our results with model predictions of different merger categories places some constraints on the possible merger progenitors. These merger events are required to happen at high-redshift (i.e., $z \geq 1$), between progenitors of different mass ratio (at least 3:1) and containing a significant amount of gas (i.e., ≥ 10 percent). A further scenario is that the low-luminosity galaxies were originally late-type galaxies, whose star formation has been truncated by removal of gas and subsequently the disc has been dynamically heated by high speed encounters in the cluster environment.

Key words: galaxies: dwarf – galaxies: elliptical and lenticular, cD – galaxies: formation – galaxies: evolution – galaxies: photometry – galaxies: kinematics and dynamics

1 INTRODUCTION

In the past years, an increasing number of observational studies have focused their effort in probing kinematic, photometric and stellar population properties at large galactocentric radii of early-type galaxies (e.g., Hau & Forbes 2006; Spolaor et al. 2009a; Proctor et al. 2009; Weijmans et al. 2009; Foster et al. 2009). This interest is motivated by the fact that nuclear regions are indicative of only ~ 15 percent of the total stellar mass and of less than 10 percent of the angular momentum of a galaxy (i.e., assuming a de Vaucouleurs $r^{1/4}$ light profile and a flat rotation profile).

In numerical simulations, the efficiency of galaxy formation mechanisms are predicted to vary with radius leav-

ing measurable changes at different galactocentric radii. The spatial distribution of these observed stellar properties, and their relationship to galaxy structural parameters, provide unique clues to constrain competing galaxy formation processes.

A specific example is the mass-metallicity gradient relation found by Spolaor et al. (2009a). The correlation shows a sharp change in slope at a dynamical mass of $\sim 3.5 \times 10^{10} M_\odot$ (i.e., $M_B \sim -19$). The metallicity gradients of more massive galaxies (i.e., high-luminosity galaxies) flatten with increasing mass, perhaps due to the progressively larger number of mergers that these galaxies undergo to assemble their mass. In the remnants, the gradients are partially regenerated if a central burst of star formation is induced by the merger. In low-mass (i.e., low-luminosity) galaxies, an early star-forming collapse with a varying star formation efficiency is required in order to explain the steepening

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of gradients with increasing mass. The depth of the galaxy gravitational potential well coupled with energy feedback via supernova-driven galactic winds are responsible for shaping the gradients.

The effect of gas dissipation, merger-induced starbursts, and energy feedback are crucial ingredients in galaxy formation models to be able to reproduce the observed stellar properties of early-type galaxies (e.g., Kawata & Gibson 2003; Pipino & Matteucci 2006a; Springel et al. 2005; Robertson et al. 2006). Specifically, the interplay and time ordering between dissipative processes and the redistribution of angular momentum could be the primary sources of the different stellar properties observed in low-luminosity and high-luminosity galaxies (e.g., Rix et al. 1999; Di Matteo et al. 2009; Hoffman et al. 2009).

The variety of kinematic and photometric features displayed by early-type galaxies, such as isophotal shapes, extended stellar discs, amount of ordered and random motion are interpreted as signatures of these mechanisms. For example, the results of the SAURON survey (e.g., Bacon et al. 2001) have prompted the division of early-type galaxies into slow and fast rotators, depending on the amount of stellar angular momentum per unit mass inside one effective radius (Emsellem et al. 2007; Krajnović et al. 2008).

The aim of this paper is to complement the stellar populations findings of Spolaor et al. (2010b) (hereafter Paper II) on galaxy formation mechanisms acting in low-luminosity early-type galaxies. We do this by exploring the kinematic and photometric properties at large galactocentric radii of the same 14 low-luminosity galaxies.

In general, properties of low-luminosity early-type galaxies have not been fully probed to large radii (e.g., Davies et al. 1983; Bender et al. 1994; Halliday et al. 2001). The low-surface brightness of these objects makes the measurement of reliable properties at large radii observationally challenging. Here, we study the galaxy kinematics by analysing the line-of-sight velocity distribution (LOSVD) as a function of galactocentric radius from high quality Gemini GMOS long-slit spectroscopic data. We derive spatially resolved radial profiles out to $\sim 1 - 3$ effective radii for rotation velocity, velocity dispersion and the Gauss-Hermite coefficients h_3 and h_4 . The higher order moments of the distribution allow us to derive symmetric and asymmetric deviations from a pure Gaussian shape. They relate to the amount of ordered motion and the shape of stellar orbits (e.g. van der Marel & Franx 1993; Bender et al. 1994; Cappellari & Emsellem 2004). The isophotal properties of a galaxy, such as surface brightness profile, ellipticity, position angle and discyness/boxiness are obtained from high-resolution HST/ACS imaging data in the g and z filters.

The paper is organised as follows. In Section 2, we describe the data sample. In Section 3, we describe observations and reduction of spectroscopic and imaging data. Section 4 the technique used to recover stellar kinematics is described and the results are presented. In Section 5, we describe the surface photometry analysis and the results are presented. In Section 6, the combined kinematic and photometric properties of the galaxies are discussed. In Section 7, we discuss our results in the context of predictions from

competing galaxy formation scenarios. In Section 8, conclusions are given.

2 DATA SAMPLE

The data sample consists of 14 low-mass early-type galaxies in the Fornax and Virgo clusters. The six galaxies in the Fornax cluster are chosen from the catalogue of Ferguson (1989), while the eight galaxies in the Virgo cluster are from the catalogue of Binggeli et al. (1985). Galaxies of low mass are selected on the basis on their central stellar velocity dispersion σ_0 and the absolute B -band total magnitude M_B , both of which are known to be independent proxies of galaxy mass. The sample uniformly covers the range $1.6 < \log(\sigma_0) < 2.15$ and $-16.5 > M_B > -19.5$. This translates to a dynamical mass range of about $10^9 < M_{\text{dyn}}/\text{M}_{\odot} < 10^{11}$, using the $\log(M_{\text{dyn}}) = 2\log(\sigma_0) + \log(r_e) + 5.0$ where r_e is the effective radius in parsecs (Cappellari et al. 2006). The main properties of the galaxy sample are summarised in Table 1.

3 OBSERVATIONS AND DATA REDUCTION

3.1 GMOS long-slit spectroscopy

The observations were made using Gemini Multi-Object Spectrograph (GMOS) mounted on the Gemini South telescope in long-slit spectroscopy mode. The six galaxies in the Fornax cluster were observed during the semester 2006B (Program ID GS-2006B-Q-74), while the eight galaxies in the Virgo cluster were observed during the semester 2008A (Program ID GS-2008A-Q-3).

Four 1800s exposure were taken for each galaxy except for FCC 335, for which two exposures of 1800s were taken. At the beginning of each exposure, the slit was centred on the nucleus and oriented along the major axis of the galaxy. The GMOS detector comprises three 2048×4608 EEV CCDs with a pixel size of $13.5 \times 13.5 \mu\text{m}^2$. We adopted the standard full frame readout, and to enhance the signal-to-noise, we selected to bin $2 \times$ spatially, yielding a spatial resolution of $0.1454 \text{ arcsec pixel}^{-1}$.

The spectrograph slit width was set to 1.0 arcsec for the 2006B run and 0.5 arcsec for the 2008A run. It was $5.5'$ in length in both cases. For the 2006B run, we used the 600 line mm^{-1} grism blazed at 5075 \AA , providing a dispersion of $0.457 \text{ \AA pixel}^{-1}$ and a wavelength coverage of $3670 - 6500 \text{ \AA}$. For the 2008A run, we used the 600 line mm^{-1} grism blazed at 5250 \AA with $0.458 \text{ \AA pixel}^{-1}$, covering the wavelength range of $3860 - 6670 \text{ \AA}$. The instrumental set up provided a spectral resolution, as measured from the full width at half maximum (FWHM) of the arc lines, of $\sim 4.21 \text{ \AA}$ (2006B run) and $\sim 2.45 \text{ \AA}$ (2008A run). During the observations, the seeing varied from 0.7-arcsec to 1.0-arcsec FWHM, which is more than adequate for the purposes of our study. Spectrophotometric and radial velocity standard stars covering a range of spectral classes were observed as calibrators. These stars belong to both the Lick/IDS stellar library (Worthey 1994) and MILES library (Sánchez-Blázquez et al. 2006). Bias frames, dome flat-fields and copper–argon (CuAr) arcs were also taken for calibration.

Galaxy (1)	Alternative Name (2)	Hubble Type (3)	r_e [arcsec] (4)	P.A. [degree] (5)	Distance [Mpc] (6)	M_B [mag] (7)	M_{dyn} $[10^9 M_{\odot}]$ (8)	Radial Range r_e (kpc) (9)
FCC 148	NGC 1375	S0	14.9	89	18.45	-17.30	7.9	2.0 (2.66)
FCC 153	IC 1963	S0	12.4	81	18.45	-17.10	3.1	3.4 (4.02)
FCC 170	NGC 1381	S0	13.7	139	18.45	-18.88	34.6	3.4 (4.76)
FCC 277	NGC 1428	E/S0	9.8	123	18.45	-17.52	9.5	2.0 (1.90)
FCC 301	ESO 358-G059	E/S0	10.4	156	18.45	-17.02	3.9	1.0 (0.58)
FCC 335	ESO 359-G002	E/S0	15.6	47	18.45	-16.93	3.5	1.0 (1.40)
VCC 575	NGC 4318	E4	7.2	65	22.08	-17.61	7.2	2.1 (1.62)
VCC 828	NGC 4387	E5	9.1	137	17.94	-18.26	10.7	2.0 (1.58)
VCC 1025	NGC 4434	E0/S0	12.9	134	22.39	-18.72	31.6	2.0 (2.80)
VCC 1146	NGC 4458	E1	26.1	175	16.37	-18.13	26.9	0.8 (1.66)
VCC 1178	NGC 4464	E3	6.6	14	15.84	-17.63	7.2	3.2 (1.62)
VCC 1297	NGC 4486B	E1	2.3	110	16.29	-16.77	16.2	3.0 (0.55)
VCC 1475	NGC 4515	E2	9.4	15	16.59	-17.84	7.7	2.0 (1.51)
VCC 1630	NGC 4551	E2	12.5	67	16.14	-18.06	1.0	2.2 (2.15)

Table 1. Galaxy properties. Col.(1): galaxy name from the catalogue of Ferguson (1989) and Binggeli et al. (1985); (2): alternative galaxy name; (3): morphological type from the HyperLEDA database; (4): B -band effective radius along the galaxy semi-major axis from the HyperLEDA database; (5): position angle of the major axis; (6): distance to the galaxy using the distance modulus derived from surface brightness fluctuations of Tonry et al. (2001), Mei et al. (2005) and Blakeslee et al. (2009); we applied the correction of Jensen et al. (2003) to the values of Tonry; (7): total B -band absolute magnitudes estimated from total apparent magnitudes given in the NASA/IPAC Extra-galactic Database and corrected for Galactic extinction; (8): dynamical mass of the galaxy using the $\log(M_{\text{dyn}}) = 2 \log(\sigma_0) + \log(r_e) + 5.0$, where r_e is the effective radius expressed in parsecs (Cappellari et al. 2006); (9): radial range in effective radius and kpc (in parenthesis) covered by our observations.

The data reduction was carried out using the Gemini/GMOS IRAF package (version 1.9.1). Each object frame is treated separately. Initial reduction of the CCD frames involves overscan correction, bias and dark subtraction, flat-field correction and cosmic ray removal. The flatfield correction is performed by means of both a quartz-halogen lamp (taken with the Gemini calibration unit, GCAL) and twilight spectra, which are normalised and divided from all the spectra. This process corrects for pixel-to-pixel sensitivity variations and for two-dimensional large-scale illumination patterns due to slit vignetting. Cosmic rays are identified and eliminated by interpolating each line of an image with a high order function. The residuals (cosmic ray hits) that are narrower than the expect instrumental line width are replaced with the fitted value. The package GSWEVELLENGTH is used to establish an accurate wavelength calibration from copper–argon (CuAr) arc images taken just before and after each exposure. Each spectrum is converted to a linear wavelength scale using roughly 130 arc lines fitted by 3rd-5th order polynomials. An accuracy of better than 0.1 Å is consistently achieved for the wavelength calibration solution along the entire spectral coverage. The spectra are also corrected for geometrical distortions along the spectral directions (S-distortion). The spectral resolution is derived as the mean of the Gaussian FWHMs measured for a number of unblended arc-lamp lines which are distributed over the whole spectral range of a wavelength-calibrated spectrum.

To obtain reliable measurements of the kinematic parameters at large galactocentric radii, it is critical to perform an accurate sky subtraction. Sky contamination becomes increasingly significant for the outer parts of the galaxy. In these regions, the level of galaxy light can be a few percent of the sky signal. This effect represents one of the main sources of systematic errors in the analysis of the kinematic

radial profiles. The generous field-of-view of GMOS and the intrinsic low luminosity of our sample allow us to accurately estimate the sky continuum from the edges of the slit, where the galaxy light is negligible. The sky level is estimated by interpolating along the outermost 10–20 arcsec at the two edges of the slit and subtracted from the whole image. A sky subtraction accurate to within one percent is achieved, guaranteeing negligible related systematics on the measured kinematic parameters.

For each galaxy, we obtained four fully reduced galaxy frames of identical exposure time. These are then co-added to form a single frame. The final galaxy frame is a two-dimensional spectrum from which we extract 1D spectra along the slit. The spatial width (i.e. the number of CCD rows binned) for each extracted 1D spectrum increases with radius to achieve a minimum signal-to-noise (S/N) per Å of ~ 35 in the spectral region of Mg_b . By adopting this criterion we have been able to obtain reliable measurements of kinematic parameters to galactocentric radii of $\sim 3r_e$.

3.2 HST/ACS imaging

The imaging observations were performed with the Advanced Camera for Surveys (ACS) mounted on the Hubble Space Telescope (HST). The eight Virgo cluster galaxies were observed by the ACS Virgo Cluster Survey (ACSVCS; Côté et al. 2004), while the six galaxies in the Fornax cluster were observed by the ACS Fornax Cluster Survey (ACSFCS; Jordán et al. 2007). Each galaxy was observed using the ACS Wide Field Channel (WFC) mode (Ford et al. 1998; Sirianni et al. 2005). The WFC consists of two 2048×4096 CCD detectors (WFC1 and WFC2) with a pixel size of $15\mu\text{m}$ and a spatial scale of $0.049'' \text{ pixel}^{-1}$. The nucleus of each galaxy was placed on the WFC1 detector at the reference

pixel (2096, 200). To avoid the gap between the WFC1 and WFC2 detectors an additional offset 5'' to 10'' was applied (depending on the brightness of the target galaxy).

Each galaxy was observed with the same observational methodology. Two images in the F475W filter (which resembles the SDSS *g* filter) for a total of 760s, with a respective exposure time of 380s to aid the removal of cosmic rays. Two images in the F850LP filter (similar to the SDSS *z* filter) for a total of 1220s, split in a pair of 760s exposures. A further image of 90s exposure in F850LP filter was taken to recover the central regions of the galaxies whose nuclei might saturate in the longer 560s exposure.

The data were downloaded from the Multimission Archive at the Space Telescope Science Institute (MAST). The images were processed by the ACS calibration pipeline (CALACS; Pavlovsky et al. 2006) using the “On the Fly Re-processing” option. An automated data reduction pipeline was developed for the ACSVCS data (Jordán et al. 2004). The data reductions procedures for the ACSFCS data have been documented in Jordán et al. (2007) and they are similar to those adopted for the ACSVCS data. Briefly, CALACS performed bias and dark subtraction, flat-fielding, and overscan removal. The five images of each galaxy were aligned by matching coordinates for a list of objects in common, and images taken with the same filter were then combined. A cosmic rays rejection procedure was applied before performing geometric distortion correction to combined images. The field-of-view of the fully reduced images is approximately 202'' \times 202'' in the shape of a rhomboid.

4 GALAXY STELLAR KINEMATICS

We use the penalised pixel fitting (pPXF) code of Cappellari & Emsellem (2004) to extract kinematic parameters of the galaxies. This is based on the method for identification of non-Gaussian line profiles, originally developed by van der Marel & Franx (1993).

In the pPXF software the line-of-sight velocity distribution (LOSVD) is parametrized as a sum of orthogonal functions in a Gauss-Hermite series, namely a Gaussian plus a series of Gauss-Hermite polynomials. A χ^2 minimising process determines the best-fitting moments of the LOSVD. The mean line-of-sight velocity, v_{los} , the velocity dispersion, σ , and the coefficients of the Gauss-Hermite polynomials of third, h_3 , and forth, h_4 , order.

The coefficients h_3 and h_4 quantify asymmetric (skewness) and symmetric (kurtosis) deviations of the LOSVD from a simple Gaussian distribution, respectively. Positive h_3 values indicate a distribution skewed towards velocities lower than the systemic velocity of the galaxy, v_{sys} . Conversely, negative values of h_3 are characteristic of a distribution skewed toward velocities greater than v_{sys} . Positive h_4 values correspond to a distribution more peaked than a Gaussian, while negative values of h_4 denote a flatter distribution.

The software works in pixel space, finding the linear combination of template stars spectra which, after being convolved with an appropriate LOSVD, best reproduces the galaxy spectrum. The advantage of this technique is that errors due to template-mismatch are reduced by using a combination of template stars. In our analysis, ten template

stars of spectral type in the range from G3 to K9 are used. The entire wavelength range covered by the galaxy spectrum is modelled, but the code allows us to mask out regions possible affected by skyline residuals or emission lines. A pixel-by-pixel χ^2 minimisation of the residuals to the best fit is carried out. An adjustable penalty term is applied to the χ^2 to bias the higher order moments coefficients h_3 and h_4 of the resulting LOSVD towards a Gaussian shape. We find that a small bias of 0.4 provides an optimum value for the spectra of our data sample.

We estimate random errors on the derived parameters using a series of Monte Carlo simulations. The entire kinematic measurement process is repeated on a large number of different realisations of our data sample. The model spectra are obtained using the corresponding best linear combination of template stars, broadened by the best-fitting LOSVDs and degraded to a range of S/N values.

4.1 Radial profiles

The radial profiles of the kinematic parameters v_{rot} , σ , h_3 , and h_4 for the 14 low-mass galaxies are presented in Fig. 1, Fig. 2, and Fig. 3. The parameters are expressed as a function of the galactocentric radius along the galaxy’s major-axis of each spatially resolved aperture, scaled with the effective radius of the galaxy’s semi-major axis. The rotation velocity v_{rot} is calculated as the difference between the measured v_{los} of an aperture and the v_{sys} of the galaxy, that is the median value of v_{los} from all apertures extracted along the slit for each galaxy. A description of the individual galaxies is given in Appendix A.

To be able to compare our results with previous studies we derive the following parameters describing the kinematic properties of the galaxies. The maximum measured rotation velocity v_{max} within a radius $r \leq r_e/2$. The central σ_0 and mean $\bar{\sigma}$ velocity dispersion, obtained by averaging the σ values estimated in all spatially resolved apertures within a radius $r \leq r_e/8$ and $r \leq r_e/2$, respectively. The mean of absolute values of h_3 , $|\overline{h_3}|$, and the mean of values of h_4 , $\overline{h_4}$, estimated from all spatially resolved apertures within a radius $r \leq r_e/2$. The parameters for each galaxy are listed in Table 2.

5 GALAXY SURFACE PHOTOMETRY

The isophotal analysis of images in both filters of each galaxy is conducted using the IRAF task ELLIPSE. The package is based on the method of Jedrzejewski (1987), and it involves the creation of a smooth elliptical galaxy model to fit to the real galaxy image (e.g., Ferrarese et al. 2006).

The technique makes the assumption that each isophote can be modelled by an ellipse whose centre, ellipticity and position angle are allowed to vary to reach the best-fitting values. The model is built up from a first guess elliptical isophote whose centre is pre-determined. The semi-major axis of each isophote is then incremented logarithmically from the previously fitted ellipse. The process is iterated until specific convergency criteria are achieved. A standard k-sigma clipping routine is used to evaluate and to flag divergent points, thus increasing the stability of the convergency process. A mask is created to exclude from the fit bright

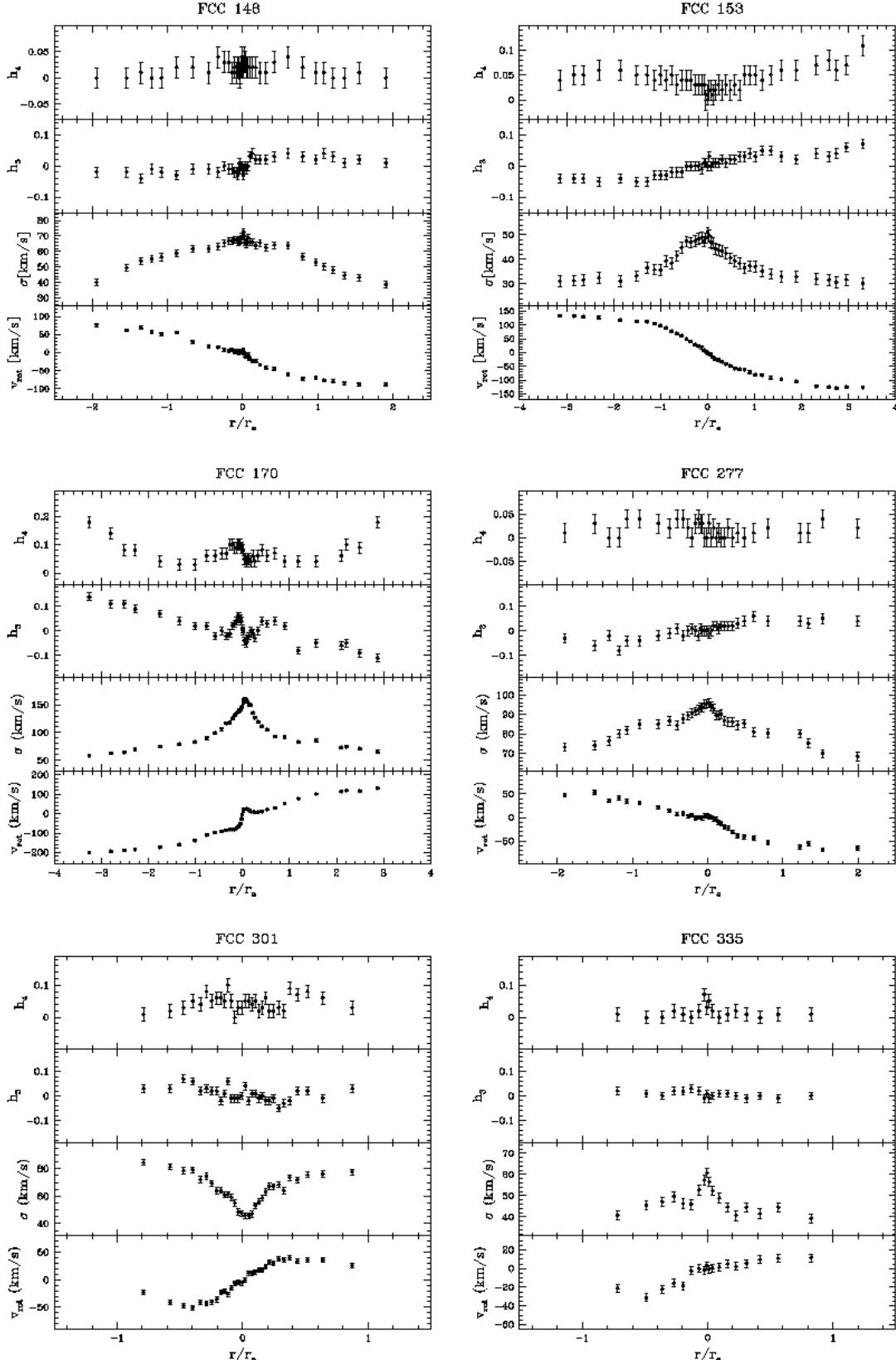


Figure 1. The major-axis stellar kinematics of the Fornax cluster low-luminosity galaxies. From bottom to top: radial profiles of rotational velocity v_{rot} , velocity dispersion σ and the Gauss-Hermite coefficients h_3 and h_4 . The kinematic quantities are expressed as a function of the galactocentric radius of each spatially resolved aperture, normalised with the effective radius of the galaxy's semi-major axis.

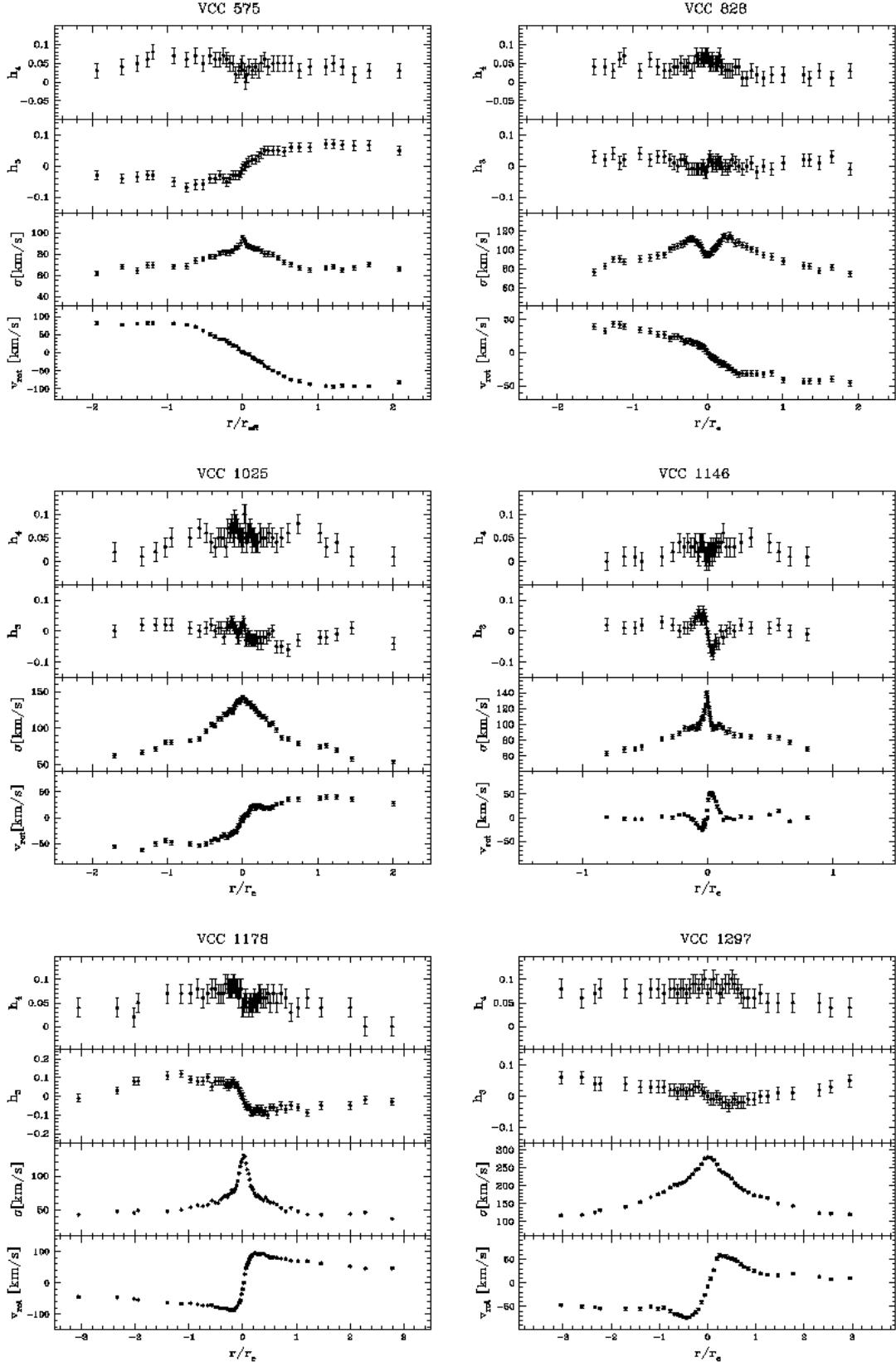


Figure 2. The major-axis stellar kinematics of the Virgo cluster low-luminosity galaxies. From bottom to top: radial profiles of rotational velocity v_{rot} , velocity dispersion σ and the Gauss-Hermite coefficients h_3 and h_4 . The kinematic quantities are expressed as a function of the galactocentric radius of each spatially resolved aperture, normalised with the effective radius of the galaxy's semi-major axis.

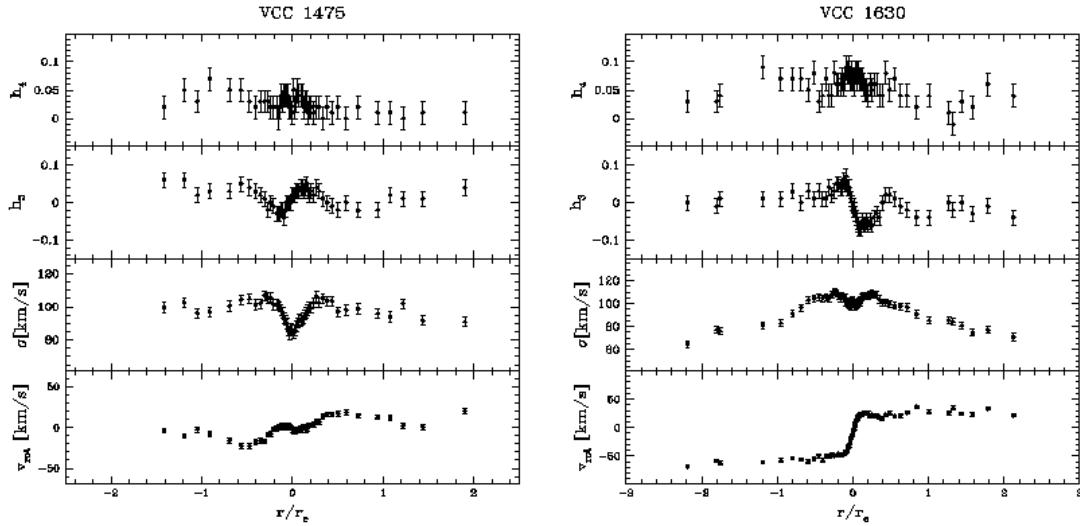


Figure 3. The major-axis stellar kinematics of the Virgo cluster low-luminosity galaxies. Continued from Fig. 2.

Galaxy	v_{sys} [km s $^{-1}$]	v_{max} [km s $^{-1}$]	σ_0 [km s $^{-1}$]	$\bar{\sigma}$ [km s $^{-1}$]	$ \bar{h}_3 $	\bar{h}_4	$\bar{g-z}$	$\bar{\epsilon}$	$P.A.$ [$^{\circ}$]	\bar{B}_4
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
FCC 148	680.1	89.0	67.8	66.1	0.031	0.021	1.04	0.589	84.2	-0.0361
FCC 153	1603.0	132.9	48.4	43.4	0.016	0.032	1.31	0.612	75.1	0.0028
FCC 170	1677.4	200.9	151.5	131.1	0.027	0.068	1.47	0.480	137.7	-0.0100
FCC 277	1595.0	68.2	93.7	89.0	0.017	0.016	1.37	0.368	116.3	-0.0079
FCC 301	1023.4	50.9	58.9	71.8	0.088	0.044	1.25	0.440	149.2	0.0004
FCC 335	1467.3	31.2	44.5	38.7	0.006	0.005	1.15	0.398	135.4	0.0014
VCC 575	1168.3	92.3	88.2	80.5	0.037	0.046	1.34	0.387	130.2	0.0071
VCC 828	543.6	46.4	104.1	102.6	0.012	0.044	1.47	0.378	36.2	-0.0244
VCC 1025	994.4	62.2	135.0	120.6	0.020	0.057	1.52	0.050	34.5	0.0014
VCC 1146	580.5	53.4	102.8	97.7	0.029	0.025	1.33	0.089	81.6	0.0034
VCC 1178	1265.6	94.7	108.2	77.7	0.070	0.065	1.32	0.302	76.2	0.0071
VCC 1297	1485.3	74.0	272.7	223.2	0.057	0.080	1.50	0.144	22.4	0.0058
VCC 1475	927.6	22.4	91.8	98.8	0.023	0.027	1.34	0.178	55.4	0.0076
VCC 1630	1187.9	68.4	109.0	109.9	0.035	0.063	1.32	0.288	148.0	-0.0107

Table 2. Kinematic and photometric parameters for the 14 low-luminosity early-type galaxies. See text for details.

globular clusters, foreground stars, background galaxies and the gap between the individual CCD detectors.

The background-subtraction is based on the average of median values from several 10×10 pixel boxes sampled in regions not affected by galaxy light. The conversion from the HST/ACS instrumental F475W and F850W filters system to the AB photometric system is performed using the synthetic coefficients from Sirianni et al. (2005). In particular, the F475W and F850W filters are transformed into g_{AB} and z_{AB} , respectively. Final magnitudes were corrected for Galactic extinction using the reddening values from the DIRBE dust maps of Schlegel et al. (1998).

The errors in magnitude are computed from the rms scatter of intensity data along the fitted ellipse and the errors in the background subtraction. Errors in the Galactic extinction correction values and in the zero-points are negligible with respect to the errors in the intensity profiles. Ellipse geometry parameter errors are obtained from the internal errors in the harmonic fit. Harmonic amplitude errors

are obtained from the fit errors after removal of all harmonics up to and including, the one being considered.

5.1 Radial profiles

We present the results of the isophotal analysis for the 14 sample galaxies in Fig. 4, Fig. 5, Fig. 6, and Fig. 7. Photometric parameters are plotted as function of the geometric mean radius r , defined as $r = a[1 - \epsilon(a)]^{1/2}$ where a is the length of the isophotal semi-major axis. This radius is then scaled with the galaxy's effective radius along the semi-major axis and expressed as r/r_e .

For each galaxy, the μ_g , μ_z surface brightness profiles and the $g - z$ colour index profile are shown. We also plot radial profiles of isophotal geometry parameters of ellipticity ϵ , position angle $P.A.$ and the isophote's deviations from perfect ellipticity A_3 , A_4 , and B_3 , B_4 . These parameters are respectively the third and fourth cosine and sine terms of a

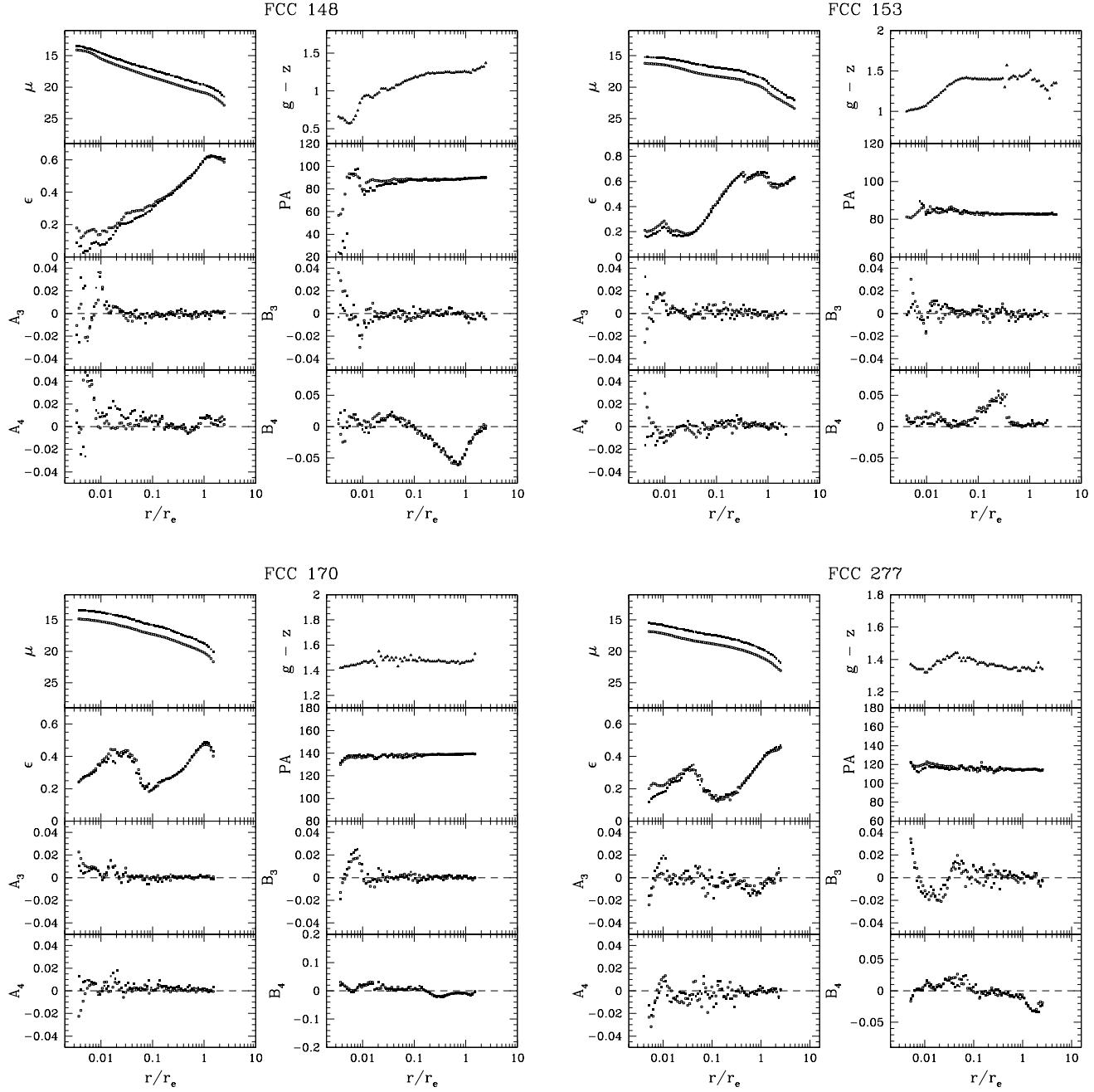


Figure 4. Isophotal properties of the Fornax cluster low-luminosity galaxies. Open black squares indicate g -band parameters. Filled black squares denote z -band parameters. In each plot are shown: surface brightness profiles μ and colour index profiles; ellipticity ϵ , position angle $P.A.$ and best fit parameters A_3 , A_4 , and B_3 , B_4 . Positive values of B_4 indicate discy isophotes, whereas negative values represent boxy isophotes. The photometric quantities are expressed as a function of the galactocentric radius of data points, normalised with the effective radius of the galaxy's semi-major axis.

Fourier expansion used by the software. The coefficients A_3 and B_3 indicate isophotes with three-fold deviations from ellipses, i.e. egg-shaped or heart-shaped. Four-fold deviations are represented by the A_4 and B_4 coefficients, such that a rhomboidal or diamond shaped isophote translates into a nonzero A_4 value. A positive value for B_4 indicates discy isophotes, whereas a negative value represents boxy isophotes.

Photometric parameters are derived by averaging all

values measured inside half the effective radius in both filters. Mean parameter values of colour index $\overline{g - z}$, ellipticity $\overline{\epsilon}$, position angle $\overline{P.A.}$, coefficient $\overline{B_4}$ for all 14 galaxies are listed in Table 2.

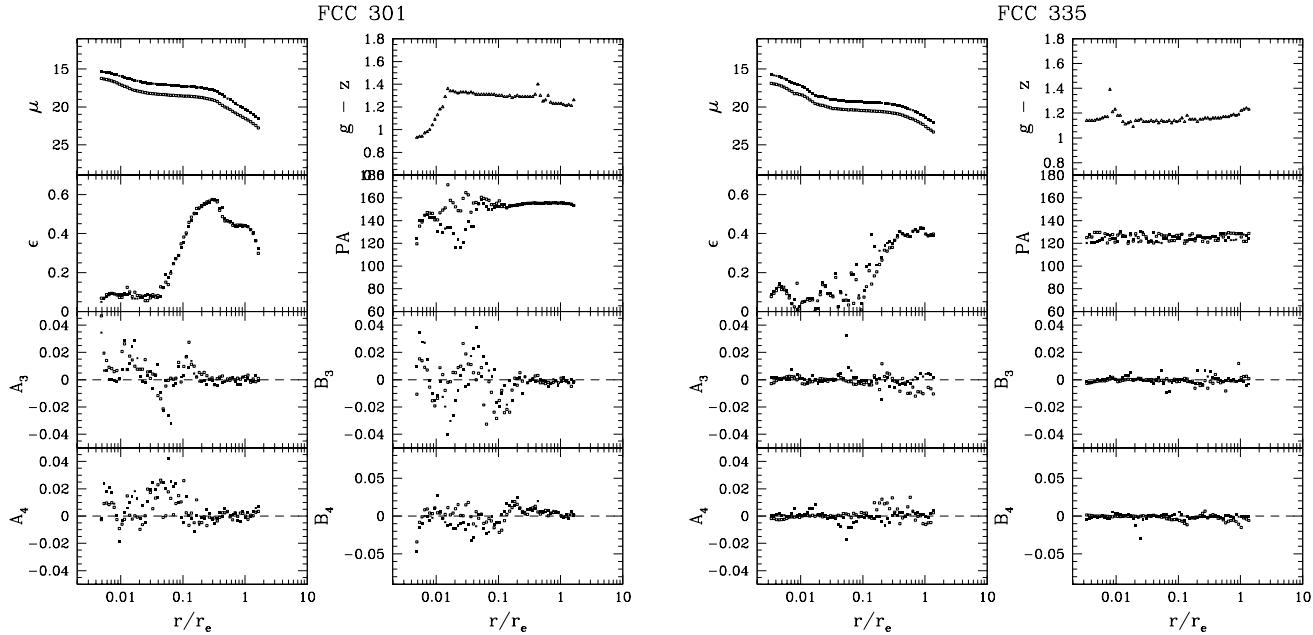


Figure 5. Isophotal properties of the Fornax cluster low-luminosity galaxies. Continued from Fig. 4.

6 GALAXY PROPERTIES

6.1 v_{rot}/σ radial profiles

The radial profiles of the ratio between the local rotation velocity and local velocity dispersion for our 14 low-luminosity galaxies are shown in Fig. 8, and Fig. 9. This ratio allows us to assess the level of ordered versus random motion as a function of galactocentric radius. The profiles have been folded onto the positive side of the major axis, with a sign change for the rotation velocity.

For each galaxy, we find that the profiles of the two major axis's sides are similar within the errors. This suggests an overall symmetry between the two sides, and thus a large-scale kinematic coherence.

In general, the profiles raise with the distance to centre of galaxy. Inside one effective radius the value of v_{rot}/σ ratio covers the range 0 - 1.4. At two effective radii the galaxies have v/σ ratio values in the range 0.6 - 2.4. This is attributed to the rotation velocity increasing with radius and contemporarily the velocity dispersion decreasing. In other words, the galaxy kinematics is primarily driven by the ordered motion of stars.

6.2 Rotational support *versus* Anisotropy

We present the $v_{max}/\bar{\sigma} - \bar{\epsilon}$ diagram (Binney 1978) for our 14 galaxies in Fig. 10. The plot allows us to quantify the amount of rotational support of these objects with respect to their intrinsic shape. We adopt as variables the ratio of v_{max} to the mean stellar velocity dispersion $\bar{\sigma}$ and the mean isophotal ellipticity $\bar{\epsilon}$.

In the diagram we also show the relationship for oblate isotropic rotator models from Binney (1978). The model predictions are derived from the tensor virial theorem, assuming that elliptical galaxies are oblate spheroids of constant ellipticity and with isotropic velocity distributions.

The values measured for our galaxies are generally consistent with the theoretical relationship, thus being oblate rotationally flattened objects. In the plot we also show the data for low-luminosity elliptical galaxies in the Virgo cluster and Leo group from Halliday et al. (2001). Four of their 14 galaxies are in common with our data sample. Their measurements are in good agreement with ours, however, we decide to use our values.

In Fig. 11, we plot the anisotropy parameter $(v/\sigma)^*$ against the total B -band absolute magnitude for each galaxy. The anisotropy parameter expresses the ratio between the observed value of $v_{max}/\bar{\sigma}$ and the $(v/\sigma)_{OI}$ value expected for an oblate rotator (Binney 1978; Davies et al. 1983). We define $(v/\sigma)^*$ in terms of observable quantities as $(v_{max}/\bar{\sigma})/\sqrt{\bar{\epsilon}/(1-\bar{\epsilon})}$ (e.g. Kormendy 1982; Bender & Nieto 1990). We also plotted the data for dwarf elliptical galaxies from Pedraz et al. (2002), Geha et al. (2003) and van Zee et al. (2004), low-luminosity ellipticals in Halliday et al. (2001) and high-luminosity elliptical galaxies from Bender et al. (1994).

The anisotropy parameter is independent of ellipticity and projection effects, thus it can be used to separate galaxies with different dynamical support (e.g., Davies et al. 1983; Scorza & Bender 1995). Objects with $(v/\sigma)^* < 0.7$ are anisotropically supported systems, whereas values of $(v/\sigma)^* \geq 0.7$ define objects primarily supported by rotation. Despite the scatter, the results suggest that high-luminosity elliptical galaxies are predominantly anisotropic objects and the amount of rotational support becomes increasingly significant towards intermediate and low-luminosity galaxies. The rotational component of the latter is significantly strong such that their observed flattened morphology is attributed to fast rotation. Dwarf elliptical galaxies tend to deviate from this trend, although exhibiting a range of rotational properties.

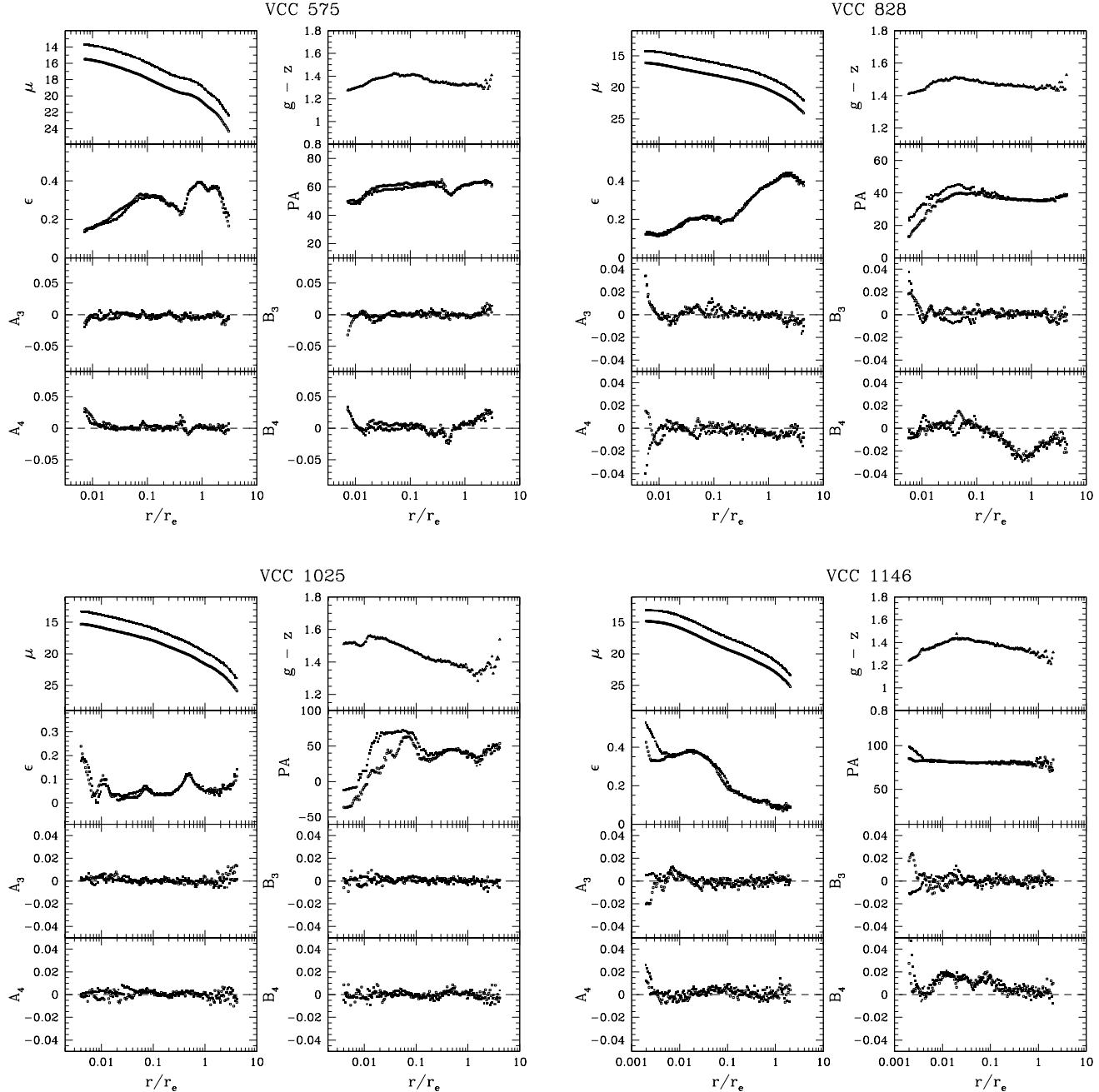


Figure 6. Isophotal properties of the Virgo cluster low-luminosity galaxies. Open black squares indicate g -band parameters. Filled black squares denote z -band parameters. In each plot are shown: surface brightness profiles μ and colour index profiles; ellipticity ϵ , position angle $P.A.$ and best fit parameters A_3 , A_4 , and B_3 , B_4 . Positive values of B_4 indicate discy isophotes, whereas negative values represent boxy isophotes. The photometric quantities are expressed as a function of the galactocentric radius of data points, normalised with the effective radius of the galaxy's semi-major axis.

6.3 Rotational support *versus* Discyness/Boxiness

In Fig. 12, the relation between amount of rotational support and isophotal shape is examined. We plot anisotropy parameter against the \overline{B}_4 coefficient of our sample galaxies, and the values of bright ellipticals by Bender et al. (1994) for comparison.

There exists a relation such that galaxies with discy isophotes tend to be more rotationally supported than boxy

galaxies. The relation is independent of galaxy's luminosity (e.g., Bender et al. 1994). We find that almost 80 percent of the low-luminosity galaxies have discy-shaped isophotes or in other terms an excess of light along the galaxy's major and/or minor axis. The remaining objects have boxy isophotes, for which the excess of light lies along a line at 45° with respect to the galaxy's axes.

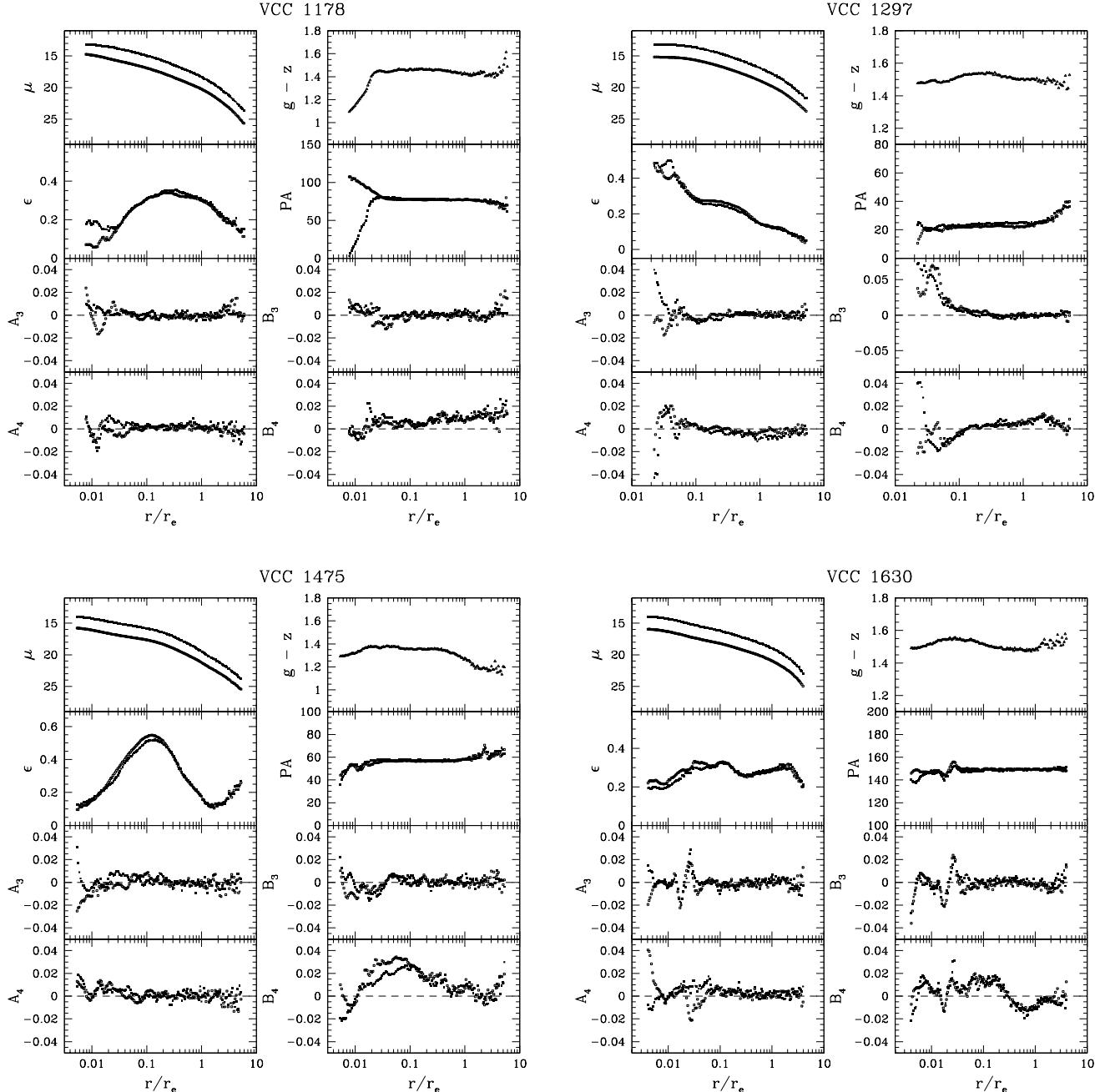


Figure 7. Isophotal properties of the Virgo cluster low-luminosity galaxies. Continued from Fig. 6.

6.4 LOSVD deviations *versus* Rotational support

The plot of the skewness coefficient h_3 against the rotational support parameter v_{rot}/σ is shown in Fig. 13. Each point corresponds to a local position along the galactocentric radius, and data for all 14 low-luminosity galaxies are presented in the same diagram.

The two parameters define a negative correlation, often referred as the $h_3 - v_{rot}/\sigma$ plane (e.g. Bender et al. 1994). In particular, in the central regions of galaxies, which corresponds to low $|v_{rot}/\sigma|$ values, the coefficient $|h_3|$ increases rapidly with $|v_{rot}/\sigma|$. At high values of $|v_{rot}/\sigma|$, characteristic of intermediate and large radii, $|h_3|$ reaches its maximum

values, up to $|h_3| \sim 0.15$. The relation was firstly observed by Bender et al. (1994), which noted that its slope is too steep to be explained by two-integral oblate rotators models (e.g., Dehnen & Gerhard 1994). Instead, they suggested that ellipticals harbouring large-scale stellar discs (~ 15 percent in mass, scalength $\sim r_e$) could explain the $h_3 - v/\sigma$ plane. Our results for low-luminosity galaxies confirm the validity at large radii $r > 1r_e$ of the $h_3 - v_{rot}/\sigma$ distribution observed by Bender et al. (1994).

We now turn to the kurtosis parameter h_4 , estimating the symmetric deviations of the LOSVDs from a Gaussian shape. In Fig. 14 we present the $h_4 - v_{rot}/\sigma$ distribution for all the low-luminosity galaxies. In contrast

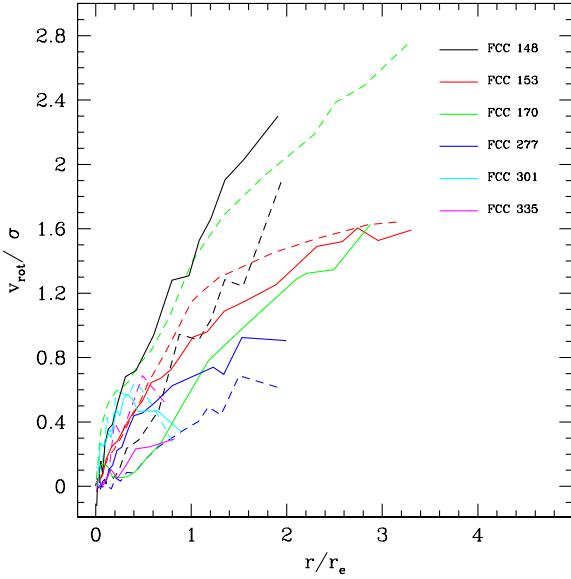


Figure 8. Radial profiles of the rotational support parameter (ratio of local rotation velocity to local velocity dispersion) for the six Fornax cluster galaxies. The quantity is expressed as a function of the ratio between the local radius and the galaxy's effective radius r/r_e . Continuous and dashed lines indicate the two sides with respect to the galaxy centre.

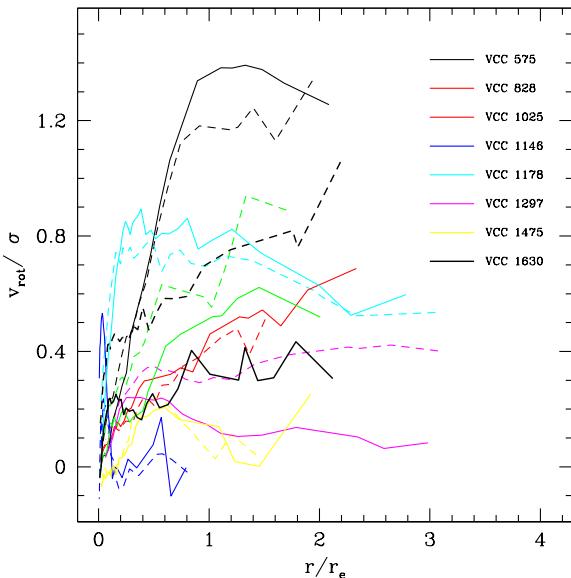


Figure 9. Radial profiles of the rotational support parameter (ratio of local rotation velocity to local velocity dispersion) for the eight Virgo cluster galaxies. The quantity is expressed as a function of the ratio between the local radius and the galaxy's effective radius r/r_e . Continuous and dashed lines indicate the two sides with respect to the galaxy centre.

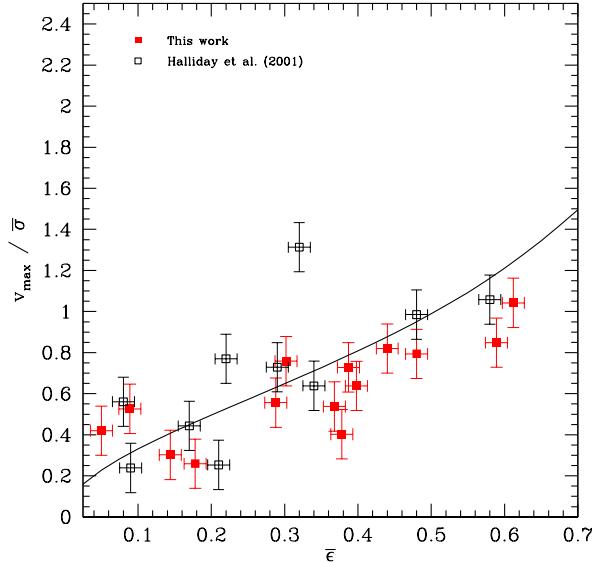


Figure 10. Ratio of maximum rotation velocity v_{max} to mean velocity dispersion $\bar{\sigma}$ plotted against mean isophotal ellipticity $\bar{\epsilon}$. Filled red squares are the data for the 14 low-luminosity galaxies presented in this paper. Open black squares are values by Halliday et al. (2001) for low-luminosity elliptical galaxies. The solid black line represents the expected relationship for an oblate, isotropic model galaxy flattened from Binney (1978).

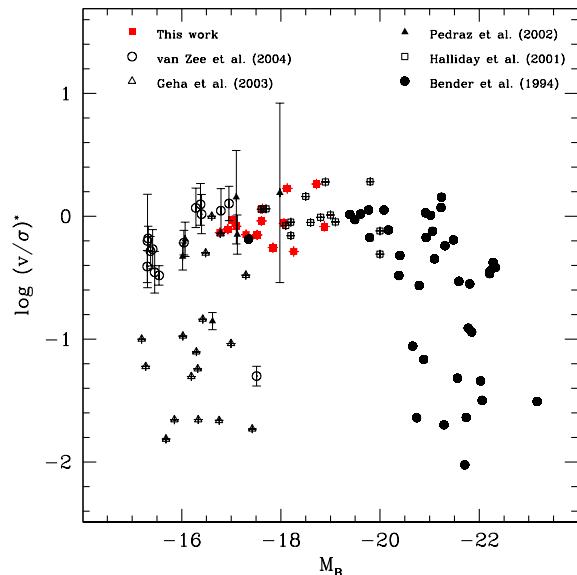


Figure 11. Logarithm of the anisotropy parameter $(v/\sigma)^*$ (ratio of the observed value of $v_{max}/\bar{\sigma}$ to the value predicted by the isotropic model) plotted against the total B -band absolute magnitude. Filled red squares are the data for the 14 low-luminosity galaxies presented in this paper. Open black squares are values by Halliday et al. (2001) for low-luminosity elliptical galaxies. Open black circles, open black triangles and filled black triangles are values for dwarf elliptical galaxies respectively from van Zee et al. (2004), Geha et al. (2003) and Pedraz et al. (2002). Filled black circles are values for bright elliptical galaxies by Bender et al. (1994).

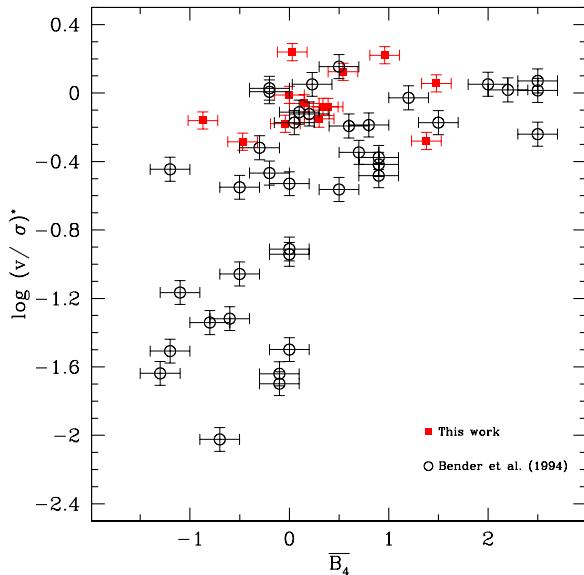


Figure 12. Anisotropy parameter against the \overline{B}_4 coefficient. Filled red squares are the data for the 14 low-luminosity galaxies presented in this paper. Open black circles are values by Bender et al. (1994) for bright elliptical galaxies. Positive values of B_4 indicate discy isophotes, whereas negative values represent boxy isophotes.

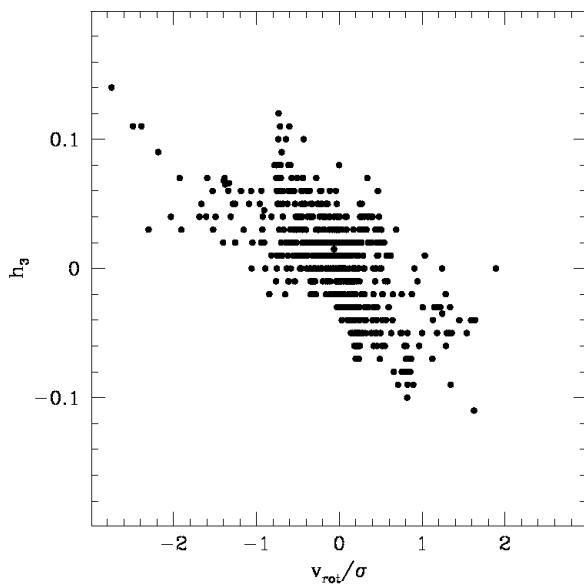


Figure 13. Local relation between the Gauss-Hermite coefficient h_3 and the rotational parameter v/σ . Data points are from the 14 low-luminosity galaxies presented in this paper.

to h_3 , h_4 is not observed to correlate with the rotational parameter. This lack of correlation was already found by Bender et al. (1994). However, the distribution present a wide spread of positive h_4 values at any v_{rot}/σ , indicating a significant rotationally-supported disc component.

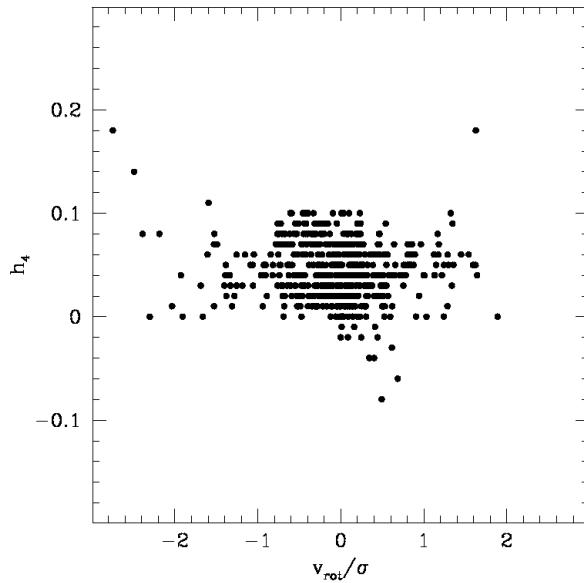


Figure 14. Local relation between the Gauss-Hermite coefficient h_4 and the rotational parameter v/σ . Data points are from the 14 low-luminosity galaxies presented in this paper.

7 DISCUSSION

We find that almost 85 per cent of our sample galaxies are characterised by a cold stellar component extending out to the outermost regions reached by our study, i.e. $\sim 3r_e$. The stellar velocity dispersion is observed to quickly decline with radius, such that ordered motion prevails over random stellar orbits. The galaxies are well described by oblate rotationally supported models. The results of our analysis confirms previous circumstantial findings by Bender & Nieto (1990), Davies et al. (1983), and Halliday et al. (2001). However, unlike these earlier studies, we are able to spatially resolved enough apertures at large galactocentric radii (i.e., $r \geq 0.5r_e$), where the data have a sufficiently high S/N to make robust measurements.

During an early star-forming collapse (e.g., Chiosi & Carraro 2002; Kawata & Gibson 2003), a cold and extended stellar component can naturally originate from the high angular momentum gas situated in the outer regions of the forming galaxies (e.g., Bekki 1998; Rix et al. 1999; Naab & Burkert 2001; Kawata & Gibson 2003). On the other hand, the gas located at inner radii tend to sink toward the galactic centre where it forms new stars. This outside-in formation scenario (e.g., Pipino et al. 2006b) is predicted to induce a shallow positive age gradient in the galactic stellar population. In Paper II we find that the 14 low-luminosity galaxies have an average age gradient of 0.06 ± 0.13 dex per decade. The gradients are such that central regions (i.e., $r \leq 0.5r_e$) are a few Gyr younger with respect to the outer parts. The stellar age of these outer galactic regions is on average older than 8 Gyr (see table 3 in Paper II). In other words, the cold stellar component of the sample galaxies is old and it was formed at high redshift, i.e. $z \geq 1$.

The radial variation in the ellipticity of the isophotes, and their prevalent discyness, observed in our data sam-

ple could be attributed to the central concentration of stars produced by the sinking gas (e.g., Mihos & Hernquist 1996; Naab & Burkert 2001). This mass congregation makes the gravitational potential well of the galaxy steeper, thus prompting a change of the stellar orbits. Specifically, it becomes more axisymmetric and less triaxial. The lack of support for box orbits makes the minor-axis tubes more populated, thus producing discy orbits.

The properties of the LOSVDs are also affected by the presence of gas (e.g., Naab et al. 2006; González-García et al. 2006). The local correlations between $h_{3,4}$ and v_{rot}/σ observed for our low-luminosity galaxies, is reproduced in numerical simulations by superimposing a cold stellar disc component in the outer regions (i.e., $\geq 1r_e$) of the spheroidal body (Naab & Burkert 2001). The transfer of angular momentum between stars and gas (e.g., Mihos & Hernquist 1996) and the redistribution of angular momentum from larger radii to inner radii produce a net increase of the rotational velocity of the galaxy.

The efficiency of gas consumption is also found to be responsible in affecting the shape of the isophotes (e.g., Bekki & Shioya 1997). Discy isophotes are more likely formed in galaxies with gradual star formation, whereas galaxies with a rapid starburst tend to have boxy isophotes. The almost solar value of the $[\alpha/\mathrm{Fe}]$ abundance ratios displayed by our galaxies (Paper II) favours this idea. The observed abundance values are a consequence of the low star formation efficiency in our galaxies, thus leading to an extended star-forming episode (≥ 1 Gyr).

In conclusion, the kinematic and photometric results of our analysis are in good agreement with the interpretation of the stellar population findings in Paper II. They suggest that an early star-forming collapse with a mass-dependent star-formation efficiency is the main mechanisms acting in the formation of these low-luminosity early-type galaxies.

Galaxy mergers may also produce some of the observed features. We now compare our results to merger simulations by Naab & Burkert (2003), Naab et al. (2006) and Cox et al. (2006).

Naab & Burkert (2003) used a large set of collisionless N -body simulations to investigate the fundamental statistical properties of merger remnants. Simulations involve dissipationless binary mergers of disc galaxies with mass ratios of 1:1, 2:1, 3:1, 4:1. Projection effects are taken into account and the bulge components of the progenitor discs are left to evolve until they coalesce before the kinematic and photometric properties of the simulated early-type galaxies are measured.

Naab et al. (2006) re-simulated the merger remnants of Naab & Burkert (2003) with an additional gas component in the progenitor disc galaxies. They considered only mergers of galaxies with mass ratios of 1:1 and 3:1. Each progenitor has 10 percent of the total mass of the stellar disc replaced by gas, but substantial heating processes prevent gas cooling and hence star formation. The gas is allowed to settle into an equilibrium state before the disc galaxies are merged.

The study of Cox et al. (2006) considers mergers of equal-mass disc galaxies. Two sets of numerical simulations are generated to compare the statistical properties of early-type galaxies formed by dissipationless and dissipational mergers. Progenitor galaxies can contain a purely exponen-

tial stellar disc with a null fraction of gas, or disc where 40 percent of the total mass is in the form of gas. Gas-rich mergers include radiative cooling of gas, star formation, feedback processes and induced growth of massive central black holes. The orbits of the merging progenitors are identical to those explored by Naab & Burkert (2003) and Naab et al. (2006), allowing a direct comparison of results.

In Fig. 15, we show the distribution of the anisotropy parameter of our sample in comparison with the predictions of Naab & Burkert (2003), Naab et al. (2006) and Cox et al. (2006). The distribution of our low-luminosity galaxies has a median value of $(v_{max}/\sigma_0)^* \sim 0.86$ and the values span the range 0.4 - 2.0. The distributions of simulated remnants have significantly different median values for the degree of rotational support with respect to our galaxy sample. However, there exists a tendency for mergers with increasingly different mass ratios to produce more isotropically supported galaxies. The predictions of Naab et al. (2006) and Cox et al. (2006) show that an additional gas component in the progenitor discs has a significant impact on the galaxy dynamical support. This dissipative component increase the amount of rotation and therefore ordered motion in the remnants. The results of a Kolmogorov–Smirnov (KS) test show that our observed galaxies and galaxies simulated by Cox et al. (2006) have a difference in their anisotropy parameter distributions significant at 4σ level. On the other hand, remnants of gas-free mergers with mass ratios of 4:1 simulated by Naab & Burkert (2003) and merger models of Naab et al. (2006) with mass ratios 3:1 with progenitors having 10 percent of gas appear similarly distributed to our sample galaxies at almost the 60 per cent level.

In Fig. 16, the observed distribution of the isophote-shape parameter \overline{B}_4 (the discyness/boxiness parameter) is compared to merger predictions. The sample galaxies have discy-shaped isophotes, with a distribution median value of $\overline{B}_4 = 0.30$. The dissipationless models of Naab & Burkert (2003) produce merger remnants with more discy isophotes when progenitors of increasingly different mass ratio are considered. On the other hand, the simulations of Naab et al. (2006) suggest that the presence of a gas component in the progenitors is responsible for changing the stellar orbits of the remnant. In this case, progenitors with a mass ratio 3:1 and 10 percent of a gas component produce remnants with a small amount of discy isophotes. A KS test marginally rejects at the 70 per cent level the hypothesis that dissipationless merger with mass ratios of 2:1 and 10 per cent gas-rich mergers with mass ratios 3:1 have the same distribution of our sample galaxies. The simulations of Cox et al. (2006) show that equal-mass progenitors with 40 percent of gas and merger-induced star formation produce remnants with significantly discy isophotes. The KS test confirms that the degree of discyness of the simulated galaxies is inconsistent with that one of the observed galaxies. We note that the inclination angle under which the galaxy is observed might affect the evaluation of the isophotes's shape (i.e. projection effects).

In conclusion, the kinematic and isophotal shape predictions from merger models examined are in less good agreement with our results. Although we suggest that merging is unlikely to be the main galaxy formation mechanism of our galaxies, comparison with the model predictions provides us

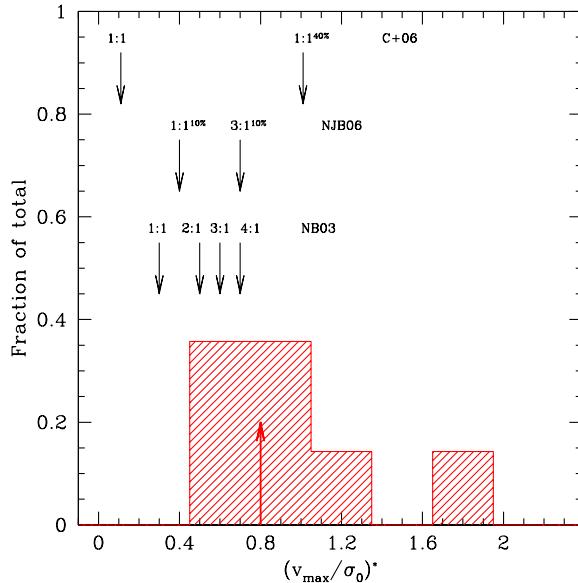


Figure 15. Frequency distributions of the anisotropy parameter for the sample of low-luminosity early-type galaxies presented in this work (shaded red region). The median value of the distribution is shown as superimposed red arrow at the bottom of the histogram. The three series of arrows above the histograms indicate the peaks of the distributions for simulated galaxies. Bottom set of arrows: simulations from Naab & Burkert (2003) (labelled as NB03) for gas-free mergers with mass ratios of 1:1, 2:1, 3:1 and 4:1. Middle set of arrows: distributions peaks from Naab et al. (2006) (labelled as NJB06) for mergers with mass ratios of 1:1 and 3:1 and progenitors with 10 percent of the total mass of the stellar disc replaced by gas. Upper set of arrows: predictions from Cox et al. (2006) (labelled as C+06) for dissipationless and dissipational equal-mass mergers; in gas-rich mergers, progenitors have 40 percent of the total mass of the stellar disc replaced by gas.

with some constraints on the merger properties. In order to reproduce the observed properties the mergers are required to happen at high-redshift ($z \gtrsim 1$), between progenitors of different mass ratio (at least 3:1) and with a significant fraction of the total mass in the form of gas (i.e., $\gtrsim 10$ percent). However, this formation scenario is not able to recreate the the observed mass-metallicity gradient relation of these low-luminosity galaxies (Spolaor et al. 2009a; Paper II).

A further interpretation of our results is that low-luminosity galaxies were originally late-type galaxies, whose star formation has been truncated by removal of gas (i.e., strangulation) and subsequently the disc has been dynamically heated by high velocity encounters (i.e., galaxy harassment) in the cluster environment. Simulations have shown that late-type galaxies entering in a rich cluster can undergo a significant morphological transformation into spheroidals by encounters with brighter galaxies and the cluster's tidal field (Moore et al. 1996; Mastropietro et al. 2005). In this scenario, we expect the original disc to be dynamically heated by the interactions such that stellar orbits acquire a significant velocity component perpendicular to the disc. The imprint of their previous morphological nature is preserved in the form of an embedded stellar disc (e.g., De Rijcke et al. 2003; Chilingarian et al. 2008). For exam-

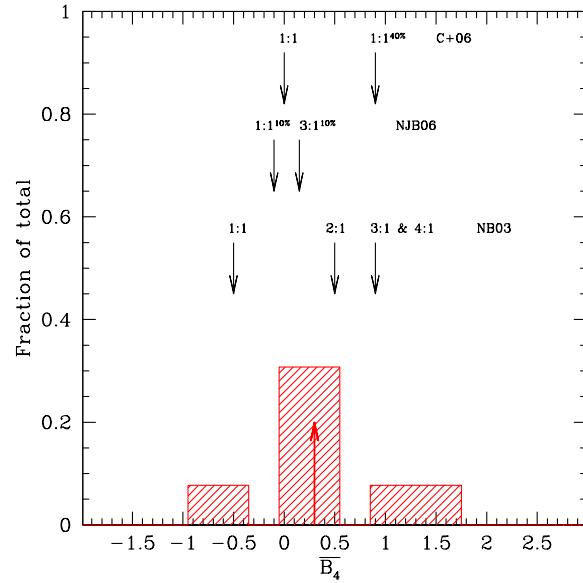


Figure 16. Frequency distributions of the coefficient \overline{B}_4 for the sample of low-luminosity early-type galaxies presented in this work (shaded red region). Description as in Fig. 15.

ple, Beasley et al. (2009) reported significant rotation at large radii (i.e., $4 - 7r_e$) of two luminous Virgo dwarf ellipticals, using globular cluster systems as tracers of galaxy dynamics. They show that the detection of such large amount of rotation in the outer galactic regions support the idea that luminous dwarf ellipticals were originally disc galaxies. Numerical simulations (Mastropietro et al. 2005) predict values of the anisotropy parameter $(v/\sigma)^*$ similar to those of our galaxies, and also disky isophotes with similar \overline{B}_4 values. Moreover, the lack of counter rotation and the high incidence of coupled rotation between disc and bulge observed in our sample galaxies may favour a disc-heating scenario, whereby stars of the original disc contribute towards the bulge population while retaining some of the angular momentum.

However, this scenario is more difficult to reconcile with the stellar population properties, and in particular the mass-metallicity gradient relation, observed in our galaxies. We expect the star formation to be truncated in the disc due to the interactions. Thus, metallicity gradients are required to form in the late-type galaxies and to be somehow preserved during the high velocity encounters. We conclude that although the kinematic and isophotal features in our galaxy sample can be interpreted in the context of a morphological transformation from late to early types, further detailed numerical simulations are needed to understand if such a scenario can also explain the stellar population trends observed.

8 CONCLUSIONS

We have investigated the kinematic and photometric properties at large galactocentric radii for a sample of 14 low-luminosity early-type galaxies. The radial extent considered in our analysis, i.e. $\sim 1 - 3r_e$, allows us to probe more than 50 percent of the total baryonic mass of a galaxy and more

than ~ 10 percent of its angular momentum. We derive radial profiles for the kinematic parameters of rotational velocity, velocity dispersion, and the Gauss-Hermite coefficients h_3 and h_4 . We also obtain radial profiles for the photometric parameters in both g and z -band of surface brightness, ellipticity, position angle and discyness/boxiness.

We find that our low-luminosity early-type galaxies host an extended, cold, and old (i.e., ≥ 8 Gyr) stellar component at radii larger than $0.5r_e$. This reaches out to the largest radii (i.e., $\sim 3r_e$) and the amount of rotational support increases as a function of the galactocentric radius. This implies that the low-luminosity galaxies are predominantly supported by rotation. The isophotes are discy-shaped and their degree of ellipticity is such that these galaxies can be well described by oblate isotropic rotator models.

The observed properties suggest that an early star-forming collapse is the main mechanism acting in the formation of these galaxies. The old and extended cold stellar component is generated from the initial high angular momentum gas in the outer regions of the forming galaxy. The low star formation efficiency of these galaxies, endorsed by their stellar population properties (Paper II), favours the observed isophotal discyness and the isophotes' ellipticity radial variation. Our results support and complement the interpretation of the stellar populations findings in Paper II.

Major mergers are unlikely to be the main mechanisms responsible for the formation of these galaxies. In particular, numerical simulations of different merger categories fail to reproduce the amount of rotational support and isophotal shape observed in the low-luminosity galaxies. Although we can not rule out a minor merging origin for these galaxies, the comparison of our results with theoretical predictions provides us with some constraints on the possible merger progenitors. In order to reproduce the observed properties, the mergers are required to happen at high-redshift (i.e., $z \geq 1$) involving progenitors of different mass ratio (at least 3:1) and with a significant amount of gas content (i.e., ≥ 10 percent).

An alternative scenario is that the low-luminosity galaxies were formerly late-type galaxies, whose star formation has been truncated by removal of gas and subsequently the disc has been dynamically heated by high speed encounters in the cluster environment. However, we conclude that further detailed numerical simulations are needed to understand if such disc-heating scenario can also reproduce the stellar population properties of our sample galaxies.

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REFERENCES

- Bacon R., et al., 2001, MNRAS, 326, 23
- Beasley M. A., Cenarro A. J., Strader J., Brodie J. P., 2009, AJ, 137, 5146
- Bekki K., Shioya Y., 1997, ApJL, 478, L17
- Bekki K., 1998, ApJL, 502, L133
- Bender R., Nieto J.-L., 1990, A&A, 239, 97
- Bender R., Saglia R. P., Gerhard O. E., 1994, MNRAS, 269, 785
- Binggeli B., Sage A., Tammann G. A., 1985, AJ, 90, 1681
- Binney J., 1978, MNRAS, 183, 501
- Blakeslee J. P., et al., 2009, ApJ, 694, 556
- Cappellari M., Emsellem E., 2004, PASP, 116, 138
- Cappellari M., et al., 2006, MNRAS, 366, 1126
- Cappellari, M., et al., 2007, MNRAS, 379, 418
- Chiosi C., Carraro G., 2002, MNRAS, 335, 335
- Chilingarian I. V., Cayatte V., Durret F., Adami C., Balkowski C., Chemin L., Laganá T. F., Prugniel P., 2008, A&A, 486, 85
- Côté P., et al., 2004, ApJS, 153, 223
- Cox T. J., Dutta S. N., Di Matteo T., Hernquist L., Hopkins P. F., Robertson B., Springel V., 2006, ApJ, 650, 791
- Cretton N., Naab T., Rix H.-W., Burkert A., 2001, ApJ, 554, 291
- Croton D. J., et al., 2006, MNRAS, 365, 11
- Davies R. L., Efstathiou G., Fall S. M., Illingworth G., Schechter P. L., 1983, ApJ, 266, 41
- De Rijcke S., Dejonghe H., Zeilinger W. W., Hau G. K. T., 2003, A&A, 400, 119
- Dehnen W., Gerhard O. E., 1994, MNRAS, 268, 1019
- Di Matteo T., Springel V., Hernquist L., 2005, Nature, 433, 604
- Di Matteo P., Jog C. J., Lehnert M. D., Combes F., Semelin B., 2009, A&A, 501, 9
- Emsellem, E., et al., 2007, MNRAS, 379, 401
- Faber S. M., et al., 1997, AJ, 114, 1771
- Ferguson H. C., 1989, AJ, 98, 367

Ferrarese L., van den Bosch F. C., Ford H. C., Jaffe W., O'Connell R. W., 1994, AJ, 108, 1598

Ferrarese L., et al., 2006, ApJS, 164, 334

Ford H. C., et al., 1998, Proc. SPIE, 3356, 234

Foster C., Proctor R. N., Forbes D. A., Spolaor M., Hopkins P. F., and Brodie J. P., 2009, MNRAS, 400, 2135

Geha M., Guhathakurta P., van der Marel R. P., 2003, AJ, 126, 1794

González-García A. C., Balcells M., Olshevsky V. S., 2006, MNRAS, 372, L78

Graham A. W., Colless M. M., Busarello G., Zaggia S., Longo G., 1998, A&AS, 133, 325

Graham A. W., Erwin P., Trujillo I., Asensio Ramos A., 2003, AJ, 125, 2951

Halliday C., Davies R. L., Kuntschner H., Birkinshaw M., Bender R., Saglia R. P., Bagley G., 2001, MNRAS, 326, 473

Hau G. K. T., Forbes D. A., 2006, MNRAS, 371, 633

Hoffman L., Cox T. J., Dutta S., and Hernquist L., 2009, ApJ, 705, 920

Hopkins P. F., Hernquist L., Cox T. J., Di Matteo T., Robertson B., Springel V., 2006, ApJS, 163, 1

Hopkins P. F., Cox T. J., Dutta S. N., Hernquist L., Kormendy J., Lauer T. R., 2009, ApJS, 181, 135

Jensen J. B., Tonry J. L., Barris B. J., Thompson R. I., Liu M. C., Rieke M. J., Ajhar E. A., Blakeslee J. P., 2003, ApJ, 583, 712

Jedrzejewski R. I., 1987, MNRAS, 226, 747

Jesseit R., Naab T., Burkert A., 2005, MNRAS, 360, 1185

Jordán A., et al., 2004, ApJS, 154, 509

Jordán A., et al., 2007, ApJS, 169, 213

Kawata D., Gibson B. K., 2003, MNRAS, 340, 908

Kormendy J., 1982, ApJ, 257, 75

Kormendy J., Bender R., 1996, ApJL, 464, L119

Krajnović D., et al., 2008, MNRAS, 390, 93

Lauer T. R., et al., 1995, AJ, 110, 2622

Mastropietro C., Moore B., Mayer L., Debattista V. P., Piffaretti R., Stadel J., 2005, MNRAS, 364, 607

Mehlert D., Saglia R. P., Bender R., Wegner G., 2000, A&AS, 141, 449

Mei S., et al., 2005, ApJ, 625, 121

Mihos J. C., Hernquist L., 1996, ApJ, 464, 641

Moore B., Katz N., Lake G., Dressler A., Oemler A., 1996, Nature, 379, 613

Naab T., Burkert A., Hernquist L., 1999, ApJL, 523, L133

Naab T., Burkert A., 2001, ApJL, 555, L91

Naab T., Burkert A., 2003, ApJ, 597, 893

Naab T., Jesseit R., Burkert A., 2006, MNRAS, 372, 839

Pavlovsky C., et al., 2006, "Advanced Camera for Surveys Instrument Hbook for Cycle 16", Version 7.1, (Baltimore: STScI).

Pedraz S., Gorgas J., Cardiel N., Sánchez-Blázquez P., Guzmán R., 2002, MNRAS, 332, L59

Pipino A. Matteucci F., 2004, MNRAS, 347, 968

Pipino A. Matteucci F., 2006, MNRAS, 365, 1114

Pipino A., Matteucci F., Chiappini C., 2006, ApJ, 638, 739

Proctor R. N., Forbes D. A., Romanowsky A. J., Brodie J. P., Strader J., Spolaor M., Mendel J. T., and Spitler L., 2009, MNRAS, 398, 91

Rix H.-W., Carollo C. M., Freeman K., 1999, ApJL, 513, L25

Robertson B., Cox T. J., Hernquist L., Franx M., Hopkins P. F., Martini P., Springel V., 2006, ApJ, 641, 21

Sánchez-Blázquez P., Gorgas J., Cardiel N., González J. J., 2006b, A&A, 457, 809

Schlegel D. J., Finkbeiner D. P., Davis M., 1998, ApJ, 500, 525

Scorza C., Bender R., 1995, A&A, 293, 20

Sirianni M., et al., 2005, PASP, 117, 1049

Spolaor M., Forbes D. A., Hau G. K. T., Proctor R. N., Brough S., 2008a, MNRAS, 385, 667

Spolaor M., Forbes D. A., Proctor R. N., Hau G. K. T., Brough S., 2008b, MNRAS, 385, 675

Spolaor M., Proctor R. N., Forbes D. A., Couch W. J., 2009a, ApJL, 691, L138

Spolaor M., Kobayashi C., Forbes D. A., Couch W. J., Hau G. K. T., 2010b, MNRAS, accepted, arXiv:1006.1698 (Paper II)

Springel V., 2000, MNRAS, 312, 859

Springel V., Di Matteo T., Hernquist L., 2005, ApJL, 620, L79

Tonry J. L., Dressler A., Blakeslee J. P., Ajhar E. A., Fletcher A. B., Luppino G. A., Metzger M. R., Moore C. B., 2001, ApJ, 546, 681

Trujillo I., Erwin P., Asensio Ramos A., Graham A. W., 2004, AJ, 127, 1917

van der Marel R. P., Franx M., 1993, ApJ, 407, 525

van Zee L., Skillman E. D., Haynes M. P., 2004, AJ, 128, 121

Weijmans A., et al., 2009, MNRAS, 398, 561

White S. D. M., 1979, MNRAS, 189, 831

Worthey G., 1994, ApJS, 95, 107

APPENDIX A: INDIVIDUAL GALAXIES

A1 FCC 148 (NGC 1375)

The rotation of FCC 148 increases outward to a maximum amplitude of $\sim 89 \text{ km s}^{-1}$, where σ falls to $\sim 40 \text{ km s}^{-1}$. The h_3 has the opposite sign to the rotation, often seen in systems with an embedded cold disk.

The increase in rotational support outwards is also reflected in the increasingly elliptical isophotes with radius, with the ellipticity rising from ~ 0.2 near the galaxy centre to > 0.6 at $r > 1r_e$ (1.3 kpc).

A2 FCC 153 (IC 1963)

The rotation of FCC 153 increases outward and plateaus at $\sim 133 \text{ km s}^{-1}$, where σ falls to $\sim 30 \text{ km s}^{-1}$. The σ profile is slightly asymmetric in the centre, being higher on the negative r side. The h_3 has the opposite sign to the rotation, often seen in systems with an embedded cold disk.

Like FCC 148 the increase in rotational support outwards is also reflected in the increasingly elliptical isophotes with radius, with the ellipticity rising from ~ 0.2 near the galaxy centre to ≈ 0.6 at $r > 0.2r_e$ (220 pc). The isophotes between 0.1 and 0.4 r_e (110 pc and 440 pc) are also significantly disky. The surface brightness μ only slightly decreases with radius in the galaxy centre, and exhibits an abrupt break at $r \sim 0.9r_e$, roughly at the transitional point where

the rotation plateaus. Thus the correlation between the surface brightness, rotation and ellipticity may be caused by a radially varying disk-to-bulge ratio.

A3 FCC 170 (NGC 1381)

The kinematics of FCC 170 exhibits several peculiarities. Firstly, over the monotonically outwardly increasing rotation, the rotation is enhanced in the central regions ($|r| < 0.3 r_e$), where σ is also peaked, and the h_3 has an opposite sign to the rotation. This seems to suggest there is a small nuclear embedded disk with radius ≈ 370 pc. In the outer parts of the galaxy the rotation is strong, at ~ 200 km s $^{-1}$, where the h_3 has an opposite sign to the rotation. This suggests substantially skewed LOSVDs at large radii, which could be the signature of an extended disk. This seems to be supported by the photometry, which shows that the isophotes are increasingly elliptical with increasing r . The ellipticity rises from ~ 0.2 at $0.1r_e$ (120 pc) to 0.5 at $1r_e$ (1.2 kpc).

A4 FCC 277 (NGC 1428)

The kinematics of FCC 277 bears some resemblance to that of FCC 148, with an outwardly rising rotation which reaches ~ 68 km s $^{-1}$. There is also a weak asymmetry in the LOSVD indicated by anti-correlation between h_3 and rotation.

The increase in rotational support outwards is also reflected in the increasingly elliptical isophotes with radius, with the ellipticity rising from ~ 0.1 at $r \sim 0.1r_e$ (90 pc) to > 0.45 at $r > 2r_e$ (1.8 kpc).

A5 FCC 301 (ESO 358-G059)

The most notable feature of FCC 301 is an outwardly rising σ profile and a depressed central σ , which is the reverse of most other galaxies. The rotation seems to peak at $|r| \sim 0.3r_e$ with amplitude of ~ 50 km s $^{-1}$ and then falls off. It is tempting to interpret the depressed central σ as due to the presence of a luminous, dynamically cold sub-component of radius $0.3r_e$ (289 pc) in the galaxy centre. Farther out, the higher σ (~ 80 km s $^{-1}$) and the lower rotation indicate an outward increase in dispersion support.

The photometry supports the hypothesis of a central kinematic subcomponent. The surface brightness is flat or slightly decreasing with radius within $0.3r_e$ (280 pc) of the galaxy centre, and falls off beyond that. The isophotes has a dramatic rise in ellipticity within the same region, reaching $\epsilon \sim 0.6$, and then falls with distance beyond that.

A6 FCC 335 (ESO 359-G002)

The rotation of FCC 335 is not symmetric. For positive r the stars' mean rotation gently rises outwards to ~ 15 km s $^{-1}$, while for negative r the stellar mean rotational amplitude reaches a maximum at ~ 30 at $0.5r_e$ (700 pc). There is also asymmetry in σ . In addition, the σ is centrally peaked inside $\sim 0.1r_e$ (140 kpc), where the h_4 is also peaked. The galaxy has an outwardly rising ellipticity, which reaches a maximum of 0.4.

A7 VCC 575 (NGC 4318)

The kinematics of VCC 575 is quite similar to that of FCC 153, but with lower rotation and σ . The presence of an embedded disk is also suggested

The increase in rotational support outwards is also reflected in the increasingly elliptical isophotes with radius, with the ellipticity rising from ~ 0.2 at $r \sim 0.4r_e$ to > 0.35 at $0.5 > r > 2r_e$. The isophotes are also significantly disky beyond $1r_e$. The surface brightness μ exhibits a break at $r \sim 0.8r_e$, roughly at the transitional point where the rotation plateaus. Thus the correlation between the surface brightness, rotation and ellipticity may be caused by a radially varying disk-to-bulge ratio.

A8 VCC 828 (NGC 4387)

The most striking feature is a distinct kinematic feature within $0.25 r_e$ (200 pc) of the nucleus, indicated by a centrally lowered dispersion and an enhanced velocity gradient compared to the rest of the galaxy. This suggests that embedded within the slowly rotating galaxy is a bright, dynamically cold small disk component with a higher rotational speed. The h_3 does not exhibit any significant feature associated with the decoupled core. This may suggest that the central disk dominates the brightness there causing only a weak deviation from a gaussian LOSVD.

The photometry also support the kinematic findings. The ellipticity profile also shows a break at the location where the dispersion changes gradient, with ϵ roughly constant at 0.2 inside $0.2 r_e$, and outward rising beyond that. There is also a break in the B_4 profile at the same location, indicating that the isophotes go from elliptical to slightly boxy from $0.2 r_e$ outwards. Our results are in good agreement with the previously kinematic measurements by Halliday et al. (2001).

A9 VCC 1025 (NGC 4434)

The kinematics of VCC 1025 shows signs of an extended embedded disk, as indicated by the asymmetric LOSVD with opposing v and h_3 . At the outerparts of the galaxy v_{rot} and σ have roughly similar amplitude, indicating comparable rotational and dispersion support. Photometrically the galaxy is fairly round throughout. It is possible that the disk is inclined.

A10 VCC 1146 (NGC 4458)

The most striking feature is the presence of a compact kinematically distinct core. The galaxy overall has little rotation and is dispersion supported. However, within $\sim 0.1r_e$ of the nucleus (210 pc), the rotation is significantly enhanced and the LOSVD is distinctly skewed, with h_3 having opposite sign to V_{rot} . Such feature can be attributed to the presence of a nuclear disk. The σ is strongly peaked at the nucleus, and falls sharply outwards. This fall is associated with the sharp rise in rotation, due to the contribution of the disk component to the galaxy surface brightness.

The compact kinematically distinct core is also detected in the photometry. There is an abrupt break in ellipticity roughly where the aforementioned kinematic transition

occurs. The nuclear region has a high ellipticity ($\epsilon \approx 0.4$) while the regions beyond $0.1r_e$ are fairly round ($\epsilon < 0.15$). The kinematically distinct core is also slightly disky compared to the rest of the galaxy. Our results are in good agreement with the previously kinematic measurements by Halliday et al. (2001).

A11 VCC 1178 (NGC 4464)

VCC 1178 appears to possess a large scale disk, as indicated by a rotation curve which resembles that of a disk galaxy, with a sharply rising rotation which reaches a maximum within $0.2r_e$ (100 pc), and then gradually declines outwards. The high asymmetry in the LOSVD also supports this hypothesis. At $r \approx 100$ pc, where the rotation reaches its maximum, the dispersion is roughly half that of the nuclear value. The kinematics is generally consistent with the presence of a dispersion-supported bulge plus a large scale disk.

The photometry also indicates a radially changing disk-to-bulge ratio. The ellipticity profile has a peak at $0.2r_e$, and is lower in the galaxy center and the outer parts. Interpreting the increase in ellipticity is caused an increase in light contribution from a disk component, both kinematics and photometry indicate that the large scale disk has a maximum contribution at $0.2 r_e$. Our results are in good agreement with the previously kinematic measurements by Halliday et al. (2001).

A12 VCC 1297 (NGC 4486B)

The kinematics of VCC 1297 bears some resemblance to that of VCC 1178, with the exception that the dispersion is comparably higher throughout the galaxy, the LOSVD asymmetry is weaker, and the decline in the dispersion profile is less dramatic. This may suggest a smaller contribution to the total surface brightness from the large-scale disk, compared to VCC 1178.

The photometry shows a generally outwardly decreasing ellipticity with radius. The ellipticity has a “shelf” between roughly 0.1 and $0.4 r_e$, where the rotational is fastest. As the rotation gradually declines to an asymptotic value, the galaxy becomes rounder. These seem to suggest that the large scale disk is more dominant at the galaxy centre than in the outer parts, and VCC 1297 has a higher support from dispersion.

A13 VCC 1475 (NGC 4515)

The most striking kinematic feature is the lowered dispersion within $0.2r_e$ (150 pc) of the nucleus. Within this region there is a mild counter-rotation of ~ 5 km s $^{-1}$ with respect to the rest of the galaxy, which has a maximum rotation of ~ 22 km s $^{-1}$. There is significant asymmetry in the LOSVD, with h_3 having the opposite sign to v_{rot} . The h_3 also changes sign on opposite sides of the nucleus. If the kinematically distinct core is caused by the light dominance of a kinematically cold subcomponent, then this subcomponent should exhibit quite pronounced counter-rotation. At first sight, it seems puzzling why the LOSVD asymmetry is pronounced, as indicated by the h_3 , while counter-rotation

is so weak as indicated by v_{rot} . It is unlikely that the kinematic subcomponent is a non-rotating dwarf or star cluster, or a disk which is orientated along the slit. It is possible that there exists a cold disk which is not aligned with the slit, such that little rotation is sampled along the slit, while there exists enough asymmetry in the LOSVD to give rise to the h_3 profile.

The photometry seems to support the above hypothesis. The P.A. is ~ 60 throughout the galaxy, thus indicating that the slit is 45 degrees offset from the major axis. The “disk” is associated with a region of high ellipticity ($\epsilon \sim 0.5$ and significant diskyness ($B_4 \sim 0.03$). Farther out, the galaxy becomes progressively round and exhibit a high degree of dispersion support.

A14 VCC 1630 (NGC 4551)

The galaxy’s rotation resembles that of a disk galaxy, with a maximum amplitude of ~ 70 km s $^{-1}$. The galaxy lies close to the dividing line between rotational and dispersion support. A striking feature is the large asymmetry in the LOSVD within $\sim 0.3r_e$ (300 pc) of the nucleus, indicated by the h_3 which has the opposite sign to v_{rot} . Over the same region σ is depressed. The kinematics are consistent with the presence of a nuclear cold disk as well as a fainter outer disk, which could be the manifestation of a single structure due to a radially varying disk-to-bulge ratio.

The photometry does not show any particular outstanding feature. The galaxy is moderately elliptical with ellipticity roughly fluctuating around $\epsilon \sim 0.3$. There is hint of a slightly enhanced ellipticity and diskyness within $r \sim 0.3r_e$. But the overall lack of a distinctive photometric feature seems to indicate the presence of a single embedded disk which give rise to the aforementioned kinematic features. Our results are in good agreement with the previously kinematic measurements by Halliday et al. (2001).